

Outline Big Biomedical/Health Data Analytic Challenges Complex-Time (kime) Spacekime Mathematical Foundation Statistical Implications of Spacekime Analytics Inferential Uncertainty Bayesian Inference Representation Applications – Longitudinal Spacekime Data Analytics Neuroimaging (UKBB, fMRI) Air quality (UCI ML Air Quality Dataset)

Big Biomedical & Health Data Analytic Challenges

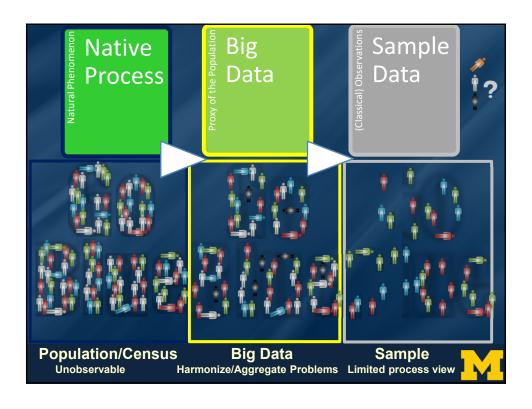


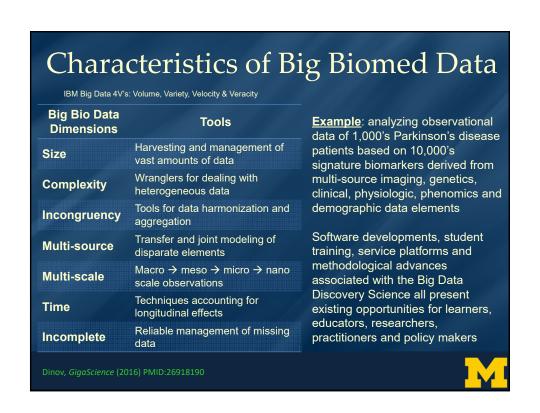
Data Analytics = Information Compression

- □ From 23 ... to ... 2^{23} (10M) $\left(\underbrace{23}_{2\#'s} \to \underbrace{2^{23}}_{8\#'s}\right)$
- Data Science: 1798 vs. 2019
- □ In the 18th century, Henry Cavendish used just 23 observations to answer a fundamental question "What is the Mass of the Earth?" He estimated very accurately the mean density of the Earth/H₂O (5.483±0.1904 g/cm³)
- □ In the 21st century to achieve the same scientific impact, matching the reliability and the precision of the Cavendish's 18th century prediction, requires a monumental community effort using massive and complex information perhaps on the order of 10M (2²³) bytes

Dinov (2016) JSMI







Data Science & Predictive Analytics

- <u>Data Science</u>: an emerging extremely transdisciplinary field bridging between the theoretical, computational, experimental, and applied areas. Deals with enormous amounts of complex, incongruent and dynamic data from multiple sources. Aims to develop algorithms, methods, tools, and services capable of ingesting such datasets and supplying semi-automated decision support systems
- ☐ Predictive Analytics: process utilizing advanced mathematical formulations, powerful statistical computing algorithms, efficient software tools, and distributed web-services to represent, interrogate, and interpret complex data. Aims to forecast trends, cluster patterns in the data, or prognosticate the process behavior either within the range or outside the range of the observed data (e.g., in the future, or at locations where data may not be available)

Data Science and Predictive Analytics

http://DSPA.predictive.space

Dinov, Springer (2018)

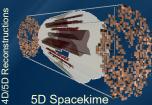


Longitudinal Data Analytics

- Neuroimaging:
 - □ 4D fMRI: time-series, represents measurements of hydrogen atom densities over a 3D lattice of spatial locations ($1 \le x, y, z \le 64$ pixels), about 3×3 millimeters² resolution. Data is recorded longitudinally over time ($1 \le t \le 180$) in increments of about 3 seconds & post-processing
 - ☐ State-of-the-art Approaches: Time-series modeling or Network analysis
 - □ Spacekime Analytics: 5D fMRI kime-series, representing the of hydrogen atom densities over the same 3D lattice of spatial locations, longitudinally over the 2D kime space, $\kappa = re^{i\varphi} \in \mathbb{C}$
 - ☐ *Differences*: Spacekime analytics estimate and utilize the kime-phases



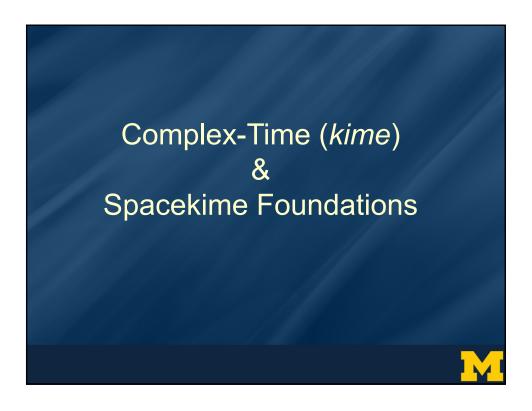


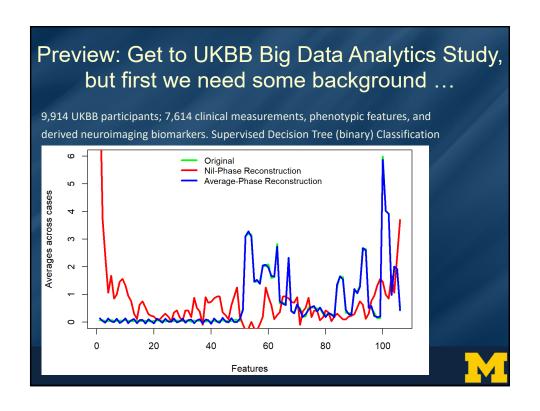


4D Spacetime

Dinov & Velev (2020)







The Fourier Transform

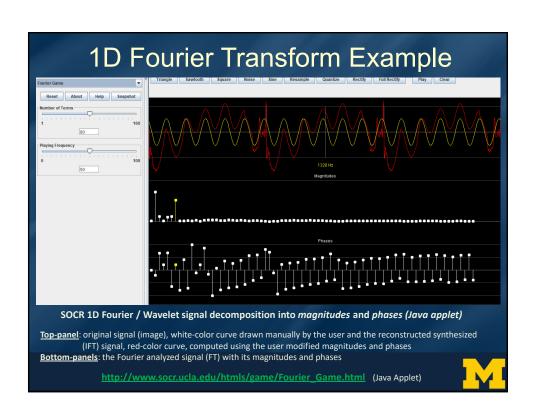
By separability, the classical **spacetime Fourier transform** is just four Fourier transforms, one for each of the four spacetime dimensions, (x,t)=(x,y,z,t). The FT is a function of the <u>angular frequency</u> ω that propagates in the wave number direction k (<u>space frequency</u>). Symbolically, the forward and inverse Fourier transforms of a 4D (n=4) spacetime function f, are defined by:

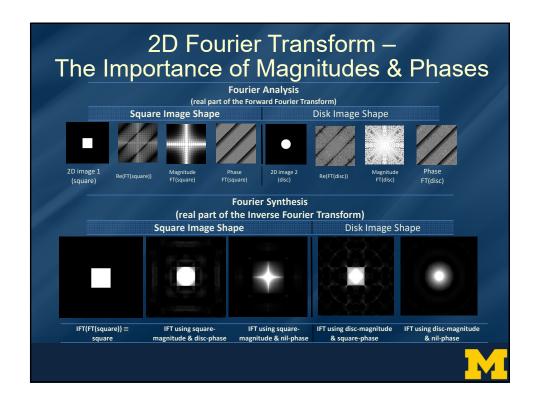
$$FT(f) = \hat{f}(\mathbf{k}, \omega) = \frac{1}{(2\pi)^{\frac{n}{2}}} \int f(\mathbf{x}, t) e^{i(\omega t - \mathbf{k}\mathbf{x})} dt d^{3}\mathbf{x},$$

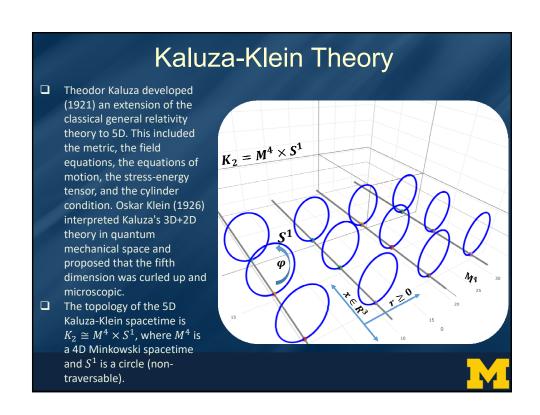
$$IFT(\hat{f}) = \hat{f}(\mathbf{x}, t) = \frac{1}{(2\pi)^{\frac{n}{2}}} \int \hat{f}(\mathbf{k}, \omega) e^{-i(\omega t - \mathbf{k}\mathbf{x})} d\omega d^{3}\mathbf{k}.$$

$$\left[\hat{f}(x,t) = IFT(\hat{f}) = IFT(FT(f)) = f(x,t)\right]$$



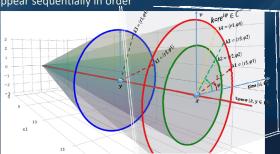






Complex-Time (Kime)

- At a given spatial location, x, complex time (kime) is defined by $\kappa = re^{i\varphi} \in \mathbb{C}$, where:
 - the magnitude represents the longitudinal events order (r > 0) and characterizes the longitudinal displacement in time, and
 - event phase $(-\pi \le \varphi < \pi)$ is an angular displacement, or event direction
- There are multiple alternative parametrizations of kime in the complex plane Space-kime universe ($\mathbb{R}^3 \times \mathbb{C}$):
 - \square (x, k1) and (x, k4) have the same spacetime representation, but different spacekime coordinates,
 - \square (x, k1) and (y, k1) share the same kime, but represent different spatial locations,
 - (x, k2) and (x, k3) have the same spatial-locations and kime-directions, but appear sequentially in order





The Spacekime Manifold

- Spacekime: $(x, k) = \left(\underbrace{x^1, x^2, x^3}_{\text{space}}, \underbrace{c\kappa_1 = x^4, c\kappa_2 = x^5}_{\text{kime}}\right) \in X, \quad c \sim 3 \times 10^8 \ m/s$
- **Kevents** (complex events): points (or states) in the spacekime manifold X. Each kevent is defined by where (x = (x, y, z)) it occurs in space, what is its causal longitudinal order $(r = \sqrt{(x^4)^2 + (x^5)^2})$, and in what kime-direction $(\varphi = \text{atan2}(x^5, x^4))$ it takes place.
- \Box Spacekime interval (ds) is defined using the general Minkowski 5×5 metric tensor $(\lambda_{ij})_{i=1,j=1}^{5,5}$, which characterizes the geometry of the *curved spacekime* manifold:

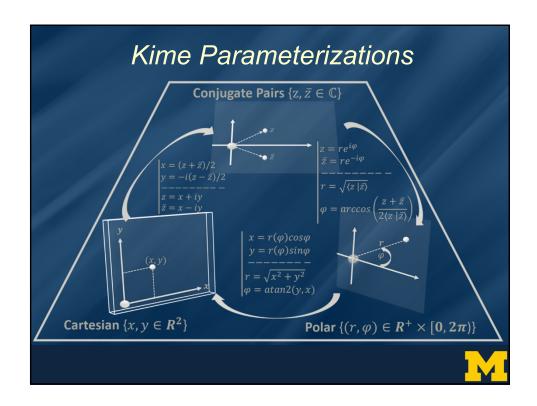
$$(\lambda_{ij})_{i=1,j=1}, \text{ which characterizes the geometry of the } \textit{curved spacekime manifold:}$$

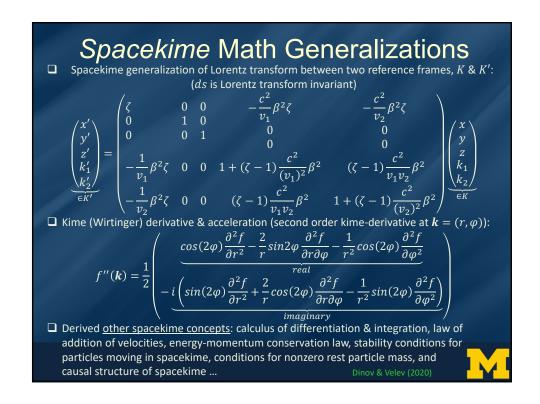
$$ds^2 = \sum_{i=1}^5 \sum_{j=1}^5 \lambda_{ij} dx^i dx^j = \lambda_{ij} dx^i dx^j$$

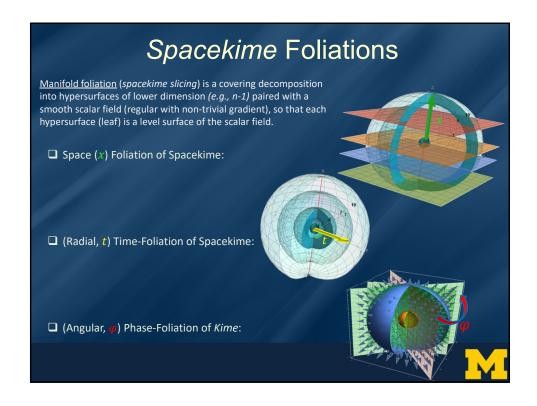
$$\text{Euclidean (flat) spacekime metric corresponds to the tensor:} \quad (\lambda_{ij}) = \begin{bmatrix} 1 & 0 & 0 & 0 & 0 \\ 0 & 1 & 0 & 0 & 0 \\ 0 & 0 & 1 & 0 & 0 \\ 0 & 0 & 0 - 1 & 0 \\ 0 & 0 & 0 & -1 & 0 \end{bmatrix}$$

- - \square Spacelike intervals correspond to $ds^2 > 0$, where an inertial frame can be found such that two kevents $a,b\in X$ are simultaneous. An object can't be present at two kevents which are separated by a spacelike interval.
 - \Box <u>Lightlike</u> intervals correspond to $ds^2 = 0$. If two kevents are on the line of a photon, then they are separated by a lightlike interval and a ray of light could travel between the two kevents.
 - \Box <u>Kimelike</u> intervals correspond to $ds^2 < 0$. An object can be present at two different kevents, which are separated by a kimelike interval.





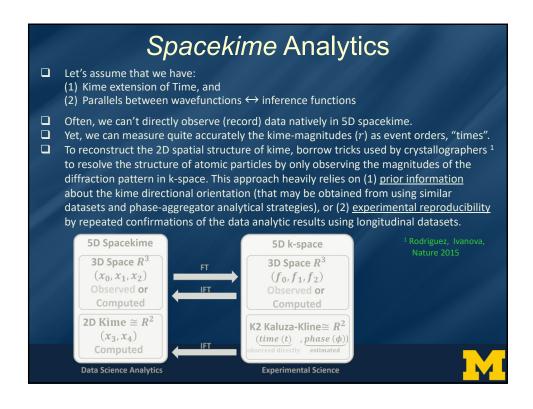


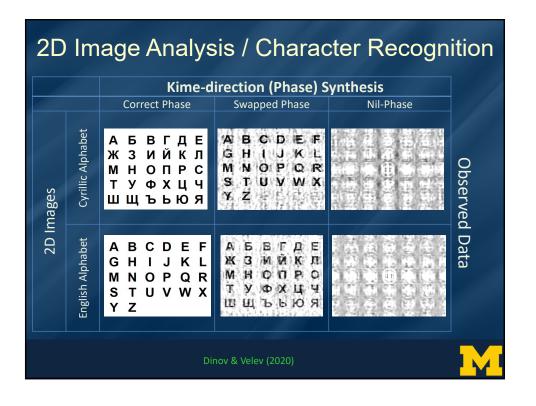


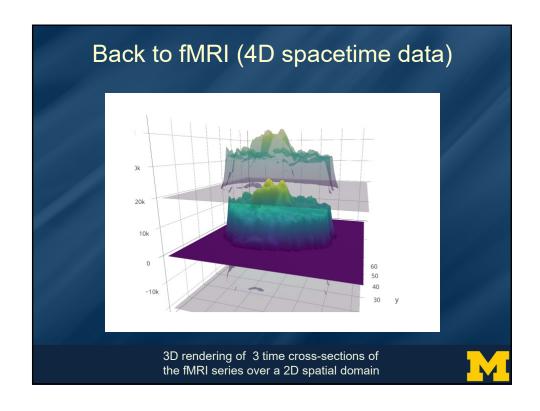


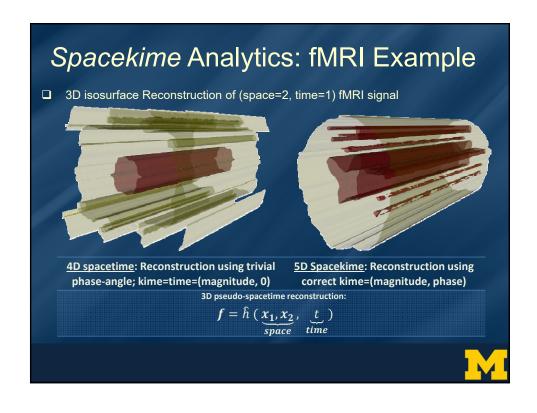
Mathematical-Physics	Data Science
A <u>particle</u> is a small localized object that permits observations and characterization of	An <u>object</u> is something that exists by itself, actually or potentially, concretely or abstractly, physically or
its physical or chemical properties An <u>observable</u> a dynamic variable about particles that can be measured	incorporeal (e.g., person, subject, etc.) A <u>feature</u> is a dynamic variable or an attribute about ar object that can be measured
Particle <u>state</u> is an observable particle characteristic (e.g., position, momentum)	<u>Datum</u> is an observed quantitative or qualitative value, an instantiation, of a feature
Particle <u>system</u> is a collection of independent particles and observable	Problem, aka Data System, is a collection of independent objects and features, without necessarily
characteristics, in a closed system Wave-function	being associated with apriori hypotheses Inference-function
Reference-Frame <u>transforms</u> (e.g., Lorentz) State of a system is an observed	Data <u>transformations</u> (e.g., wrangling, log-transform) <u>Dataset (data)</u> is an observed instance of a set of
measurement of all particles ~ wavefunction A <u>particle system is computable</u> if (1) the entire system is logical, consistent, complete	datum elements about the problem system, $0 = \{X, Y\}$ Computable data object is a very special representation of a dataset which allows direct
and (2) the unknown internal states of the system don't influence the computation	application of a dataset which allows direct application of computational processing, modeling, analytics, or inference based on the observed dataset
(wavefunction, intervals, probabilities, etc.)	analytics, or inference based on the observed dataset

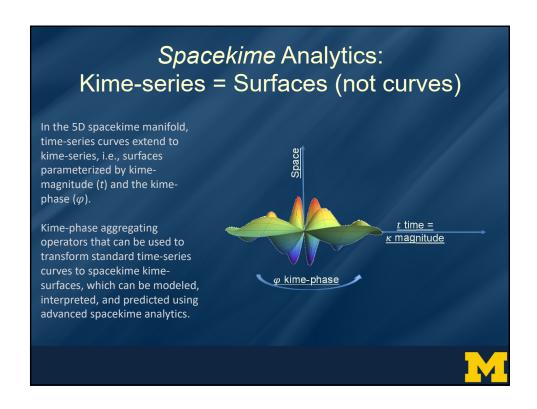
Math-Physics	Data Science				
<u>Wavefunction</u>	 Inference function - describing a solution to a specific data analytic system (a problem). For example, A linear (GLM) model represents a solution of a prediction inference problem, Y = Xβ, where the inference function quantifies the effects of all 				
Wave equ problem:	independent features (X) on the dependent outcome (Y), data: $0 = \{X,Y\}$ $\psi(0) = \psi(X,Y) = \hat{\beta} = \hat{\beta}^{OLS} = \langle X X\rangle^{-1}\langle X Y\rangle = \left(X^TX\right)^{-1}X^TY.$				
$\begin{split} &\left(\frac{\partial^2}{\partial x^2} - \frac{1}{v^2} \frac{\partial^2}{\partial t}\right) \psi(x,t) \\ &= 0 \end{split}$ Complex Solution: $\psi(x,t) = A e^{i(kx - wt)}$	$\eta \ll d$, the kernel $\psi_x(y) = \langle x y \rangle$: $0 \times 0 \to R$ transformed non-linear to linear separation, the observed data $0_i = \{x_i, y_i\} \in R^\eta$ are lifted to $\psi_{0_i} \in H$. Then, the SVM prediction operator is the weighted sum of the kernel functions at ψ_{0_i} , where β^* is a solution to the SVM regularized				
where $\left \frac{w}{k}\right = \nu$,	$\langle \psi_{O} eta^* angle_{H} = \sum_{i=1}^{n} p_i^* raket{\psi_{O} \psi_{O_i}}_{H} .$				
represents a	The linear coefficients, p_i^* , are the dual weights that are multiplied by the label corresponding to each training instance, $\{y_i\}$.				
traveling wave	Inference always depends on the (input) data; however, it does not have 1-1 and onto bijective correspondence with the data, since the inference function quantifies predictions in a probabilistic sense.				

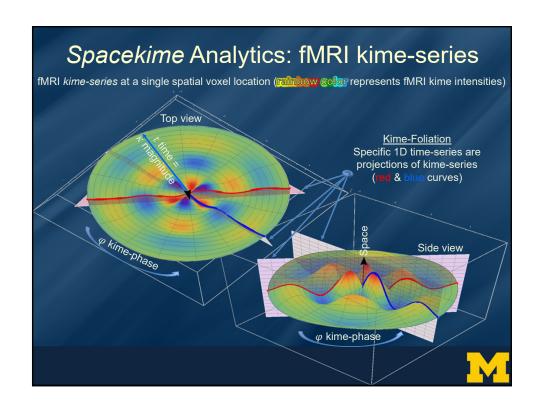


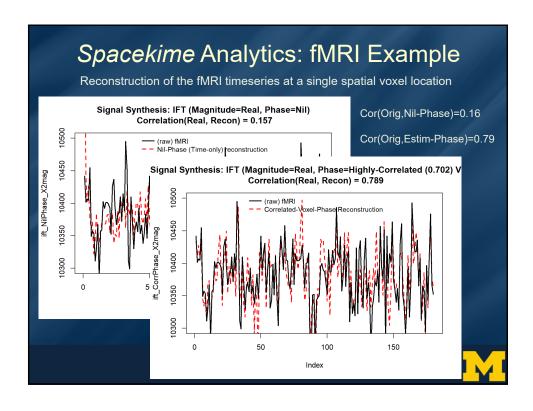














Uncertainty

- Quantum Mechanics: $||D_x u|| ||xu|| = \langle \frac{\hbar}{i} \partial_x u | ixu \rangle = \frac{\hbar}{2} ||u||^2 > 0$, i.e., non-commutation of the unbounded *operators* $D_x = \frac{\hbar}{i} \partial_x$ and x, (multiplication by x).
- Signal processing: Functions can't be time-limited and band-limited. Alternatively, a function and its Fourier transform cannot both have bounded domains $\sigma_t \times \sigma_\omega \geq 1/(4\pi)$, where σ_t, σ_ω are the time and frequency SDs.
- \Box Entropic uncertainty: Entropy can be used just like the SD to quantify distribution structure. For instance, for angular, bimodal, or divergent-variance distributions, Entropy may be a better measure of dispersion than SD. For $FT(f)(\omega) = \hat{f}(\omega)$ and $IFT(\hat{f})(x) = \hat{f}(x)$, the Shannon information entropies:

$$H_x = \int \hat{f}(x) \log \left(\hat{f}(x)\right) dx$$
 and $H_\omega = \int \hat{f}(\omega) \log \left(\hat{f}(\omega)\right) d\omega$.

satisfy: $H_x + H_\omega \ge \log(e/2)$.

L²($\mathbb R$) <u>uncertainty</u>: it is impossible for $f\in L^2$ and $\hat f$ to both decrease extremely rapidly. If both have rapidly decreasing tails: $|f(x)| \le C(1+|x|)^n e^{-a\pi x^2}$ and $|\hat f(\omega)| \le C(1+|\omega|)^n e^{-b\pi\omega^2}$, for some constant C, polynomial power n, and $a,b\in\mathbb R$, then f=0 (when ab>1); $f(x)=P_k(x)e^{-a\pi x^2}$ and $\hat f(\omega)=\widehat{P_k}(\omega)*e^{-\omega^2/4\pi a}$, where $\deg(P_k)\le n$ (when ab=1); or (when ab<1).



Heisenberg's Uncertainty in Spacekime?

- ☐ Heisenberg's uncertainty is resolved in 5D spacekime
- ☐ We can derive the classical 4D spacetime Heisenberg uncertainty as a reduction of Einstein-like 5D deterministic dynamics:
 - ☐ The math is terse it involves deriving the equations of motion by maximizing the distance (integral along the geodesic) between two points in 5D spacekime
 - The inner product $du^{\mu}\,dx_{\mu}=\frac{dx^{\mu}dx_{\mu}}{L}=\frac{ds^{2}}{L}$. Since $\frac{ds}{L}\to 1$ near the leaf membrane, $du^{\mu}\,dx_{\mu}=L=\frac{h}{mc}$. Replacing the change in velocity (du^{μ}) by the change in momentum (dp^{μ}) yields: $dp^{\mu}\,dx_{\mu}=h$.
 - This relation is similar to the quantum mechanics uncertainty principle in 4D Minkowski spacetime; however, it is obtained from 5D Einstein deterministic dynamics. In other words, in spacetime, Heisenberg's uncertainty principal manifests simply because of the one degree of freedom (kime-phase), i.e., lack of sufficient information about the second kime dimension.
 - □ In 5D spacekime, the conventional geodesic motion is perturbed by an extra force f^{μ} that can be split into two parts $f^{\mu} = f^{\mu}_{\perp} + f^{\mu}_{\parallel}$. The normal component f^{μ}_{\perp} is similar to other conventional forces and obeys the usual orthogonality condition $f^{\mu}_{\perp} u^{\mu} = 0$. However, the parallel component f^{μ}_{\parallel} has no analog in 4D spacetime. In general, it has a non-trivial inner product with the 4-velocity u^{μ} , $f^{\mu}_{\parallel} u^{\mu} \neq 0$.
- In Minkowski 4D spacetime, the <u>lack of kime-phase data naturally leaves one degree of freedom in the system causing Heisenberg's uncertainty</u>. However, the latter can be explicated by information knowledge of the fifth component (kime-phase).



Bayesian Inference Representation

- □ Suppose we have a single spacetime observation $X = \{x_{i_o}\} \sim p(x \mid \gamma)$ and $\gamma \sim p(\gamma \mid \varphi = \text{phase})$ is a process parameter (or vector) that we are trying to estimate.
- \square Spacekime analytics aims to make appropriate inference about the process X.
- □ The <u>sampling distribution</u>, $p(x \mid \gamma)$, is the distribution of the observed data X conditional on the parameter γ and the <u>prior distribution</u>, $p(\gamma \mid \varphi)$, of the parameter γ before the data X is observed, $\varphi = \text{phase aggregator}$.
- \square Assume that the hyperparameter (vector) φ , which represents the kime-phase estimates for the process, can be estimated by $\hat{\varphi} = \varphi'$.
- □ Such estimates may be obtained from an oracle, approximated using similar datasets, acquired as phases from samples of analogous processes, or derived via some phase-aggregation strategy.
- Let the <u>posterior distribution</u> of the parameter γ given the observed data $X = \{x_{i_o}\}$ be $p(\gamma|X, \varphi')$ and the process parameter distribution of the kime-phase hyperparameter vector φ be $\gamma \sim p(\gamma \mid \varphi)$.



Bayesian Inference Representation

☐ We can formulate spacekime inference as a Bayesian parameter estimation problem:

$$\underbrace{p(\gamma|X,\varphi')}_{\text{posterior distribution}} = \frac{p(\gamma,X,\varphi')}{p(X,\varphi')} = \frac{p(X|\gamma,\varphi') \times p(\gamma,\varphi')}{p(X,\varphi')} = \frac{p(X|\gamma,\varphi') \times p(\gamma,\varphi')}{p(X|\varphi') \times p(\varphi')} = \\ = \frac{p(X|\gamma,\varphi')}{p(X|\varphi')} \times \frac{p(\gamma,\varphi')}{p(\varphi')} = \underbrace{\frac{p(X|\gamma,\varphi') \times p(\gamma|\varphi')}{p(X|\varphi')}}_{\text{observed evidence}} \propto \underbrace{\frac{p(X|\gamma,\varphi')}{p(X|\varphi')} \times \frac{p(\gamma|\varphi')}{prior}}_{\text{likelihood}}.$$

- \square In Bayesian terms, the posterior probability distribution of the unknown parameter γ is proportional to the product of the likelihood and the prior.
- \square In probability terms, the posterior = likelihood times prior, divided by the observed evidence, in this case, a single spacetime data point, x_i .



Bayesian Inference Representation

- \square Spacekime analytics based on a single spacetime observation x_{i_o} can be thought of as a type of Bayesian prior-predictive *or* posterior-predictive distribution estimation problem.
- □ Prior predictive distribution of a new data point x_{j_o} , marginalized over the prior i.e., the sampling distribution $p(x_{i_o}|\gamma)$, weight-averaged by the pure prior distribution):

$$p(x_{j_o}|\varphi') = \int p(x_{j_o}|\gamma) \times \underbrace{p(\gamma|\varphi')}_{\text{prior distribution}} d\gamma.$$

□ Posterior predictive distribution of a new data point x_{j_o} , marginalized over the posterior; i.e., the sampling distribution $p(x_{j_o}|\gamma)$, weight-averaged by the posterior distribution:

$$p(x_{j_o}|x_{i_o},\varphi') = \int p(x_{j_o}|\gamma) \times \underbrace{p(\gamma|x_{i_o},\varphi')}_{\text{posterior distribution}} d\gamma.$$

- $\hfill\Box$ The difference between these $\underline{two\ predictive\ distributions}}$ is that
 - \square the posterior predictive distribution is updated by the observation $X = \{x_{i_0}\}$ and the hyperparameter, φ (phase aggregator),
 - whereas the prior predictive distribution only relies on the values of the hyperparameters that appear in the prior distribution.



Bayesian Inference Representation

- ☐ The posterior predictive distribution may be used to sample or forecast the distribution of a prospective, yet unobserved, data point x_{j_o} .
- The posterior predictive distribution spans the entire parameter statespace (Domain(γ)), just like the wavefunction represents the distribution of particle positions over the complete particle state-space.
- □ Using maximum likelihood or maximum a posteriori estimation, we can also estimate an individual parameter point-estimate, γ_o . In this frequentist approach, the point estimate may be plugged into the formula for the distribution of a data point, $p(x \mid \gamma_o)$, which enables drawing IID samples or individual outcome values.



Bayesian Inference Simulation

- \square Simulation example using 2 random samples drawn from mixture distributions each of $n_A = n_B = 10 \text{K}$ observations:
 - \square $\{X_{A,i}\}_{i=1}^{n_A}$, where $X_{A,i} = 0.3U_i + 0.7V_i$, $U_i \sim N(0,1)$ and $V_i \sim N(5,3)$, and \square $\{X_{B,i}\}_{i=1}^{n_B}$, where $X_{B,i} = 0.4P_i + 0.6Q_i$, $P_i \sim N(20,20)$ and $Q_i \sim N(100,30)$.
- □ The intensities of cohorts A and B are independent and follow different mixture distributions. We'll split the first cohort (A) into training (C) and testing (D) subgroups, and then:
 - ☐ Transform all four cohorts into Fourier k-space,
 - \square Iteratively randomly sample single observations from cohort \mathcal{C} ,
 - \square Reconstruct the data into spacetime using a single kime-magnitude value and alternative kime-phase estimates derived from cohorts B, C, and D, and
 - □ Compute the classical spacetime-derived population characteristics of cohort *A* and compare them to their spacekime counterparts obtained using a single *C* kime-magnitude paired with *B*, *C*, or *D* kime-phases.



Bayesian Inference Simulation

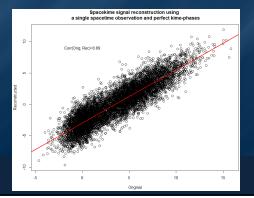
Summary statistics for the original process (cohort A) and the corresponding values of their counterparts computed using the spacekime reconstructed signals based on kime-phases of cohorts B, C, and D. The <u>estimates for the latter three cohorts correspond to reconstructions using a single spacetime observation (i.e., single kime-magnitude) and alternative kime-phases (in this case, kime-phases derived from cohorts B, C, and D).</u>

		Spacetime	Spacekime Reco	nstructions (single	kime-magnitude)
Com	nmaries	(A)	(<i>C</i>)	(B)	(D)
Sur	nmaries	Original	Phase=True	Phase=Diff. Process	Phase=Independent
	Min	-2.38798	-2.98116	-3.798440	-2.69808
1st	Quartile	-0.89359	-0.76765	-0.636799	-0.76453
	Median	0.03311	-0.05982	0.009279	-0.08329
Mean		0.00000	0.00000	0.000000	0.00000
3 rd Quartile		0.75772	0.72795	0.645119	0.69889
Max		3.61346	3.64800	3.986702	3.22987
SI	kewness	0.348269	0.2372526	0.001021943	0.31398
	Kurtosis	-0.68176	-0.4452207	0.2149918	-0.3270084
cohort xA					

Bayesian Inference Simulation

The correlation between the original data (A) and its <u>reconstruction using a single kime magnitude</u> and the correct kime-phases (C) is $\rho(A,C) = 0.89$.

This strong correlation suggests that a substantial part of the A process <u>energy</u> <u>can be recovered using only a single observation</u>. In this case, to reconstruct the signal back into spacetime and compute the corresponding correlation, we used a single kime-magnitude (sample-size=1) and process \mathcal{C} kime-phases.





Bayesian Inference Simulation

Let's demonstrate the Bayesian inference corresponding to this spacekime data analytic problem using a simulated bimodal experiment:

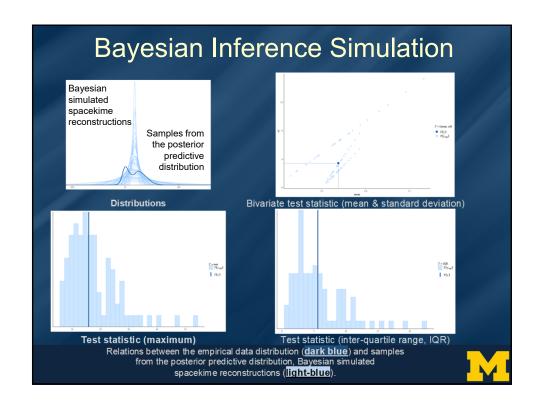
 $X_A = 0.3U + 0.7V$, where $U \sim N(0,1)$ and $V \sim N(5,3)$

Specifically, we will illustrate the Bayesian inference using repeated single spacetime observations from cohort $A, X = \{x_{i_o}\}$, and varying kime-phase priors $(\theta = \text{phase aggregator})$ obtained from cohorts B, C, or D, using different posterior predictive distributions

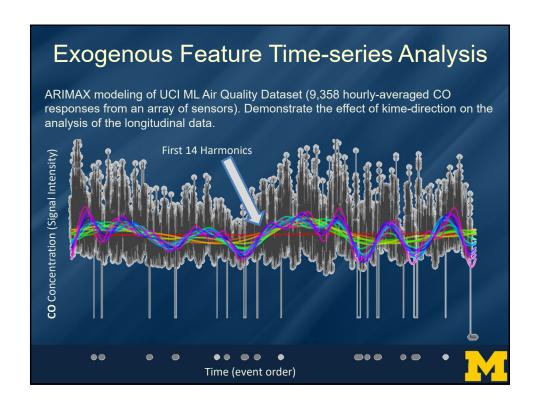
Relations between the empirical data distribution (<u>dark blue</u>) and samples from the posterior predictive distribution, representing Bayesian simulated spacekime reconstructions (<u>light-blue</u>). The derived Bayesian estimates do not perfectly match the empirical distribution of the simulated data, yet there is clearly information encoding that is captured by the spacekime data reconstructions

This <u>signal compression</u> can be exploited by subsequent model-based or model-free data analytic strategies for retrospective prediction, prospective forecasting, ML classification, derived clustering, and other spacekime inference methods









Phase	Nil	Average	True=original
Model Estimate	ARIMA(2,0,1)	ARIMA(2,0,3)	ARIMA(1,1,4)
AIC	13179	14183	10581
ar1	1.11406562	0.329482302	0.2765312
ar2	-0.14565048	0.238363531	
ma1	-0.78919188	0.267291585	-0.88913497
ma2		-0.006079386	0.12679494
ma3		0.15726556	0.03043726
ma4			-0.17655728
intercept	503.3455144	742.800113	
xreg1	-0.40283891	0.58379483	0.08035744
xreg2	0.13656613	0.280936931	6.14947902
xreg3	-0.51457636	-0.649722755	0.09859223
xreg4	1.09611981	1.239910298	0.01634736
xreg5	1.21946209	-0.026110332	-0.04816591
xreg6	1.30628469	1.081777956	-0.01104142
xreg7	1.20868397	0.254018471	0.1832854
xreg8	1.14905809	0.306524131	0.17648482
xreg9	-0.48233756	-0.405204908	6.53739782
xreg10	0.03145281	0.351063312	1.79388326
xreg11	-0.46395772	-0.457689796	-12.06965578

